

A SUBMILLIMETER SEARCH OF NEARBY YOUNG STARS FOR COLD DUST: DISCOVERY OF DEBRIS DISKS AROUND TWO LOW-MASS STARS

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ABSTRACT

We present results from a James Clerk Maxwell Telescope/SCUBA 850 μm search for cold dust around eight nearby young stars belonging to the β Pic ($t \approx 12$ Myr) and the Local Association ($t \approx 50$ Myr) moving groups. Unlike most past submillimeter studies, our sample was chosen solely on the basis of stellar age. Our observations achieve about an order of magnitude greater sensitivity in dust mass compared to previous work in this age range. We detected two of the three M dwarfs in our sample at 850 μm , GJ 182 and GJ 803 ($M_* \approx 0.5 M_\odot$), with inferred dust masses of only ≈ 0.01 – $0.03 M_\oplus$. GJ 182 may also possess a 25 μm excess, which is indicative of warm dust in the inner few AU of its disk. For GJ 803 (AU Mic; HD 197481), submillimeter mapping finds that the 850 μm emission is unresolved. A nondetection of the CO 3–2 line indicates the system is gas-poor, and the spectral energy distribution suggests the presence of a large inner disk hole (≈ 17 AU = $1''.7$ in radius for blackbody grains). These are possible indications that planets at large separations can form around M dwarfs within ~ 10 Myr. In a companion paper, we confirm the existence of a dust disk around GJ 803 using optical coronagraphic imaging. Given its youthfulness, proximity, and detectability, the GJ 803 disk will be a valuable system for studying disk, and perhaps planet, formation in great detail. Overall, submillimeter measurements of debris disks point to a drop in dust mass by a factor of $\sim 10^3$ within the first ~ 10 Myr, with the subsequent decline in the masses of submillimeter-detected disks consistent with $t^{-0.5}$ – t^{-1} .

Subject headings: circumstellar matter — planetary systems: formation —
planetary systems: protoplanetary disks — stars: late-type

1. INTRODUCTION

Circumstellar disks are common around stars at ages of a few million years. Rotationally supported disks are the natural outcome of the star formation process because of the need to conserve angular momentum during the collapse of the natal molecular core. In addition, disks are central to current theories of planet formation. These young disks are composed of gas and dust, with estimated masses as large as several percent of the stellar mass. By virtue of their large optical depths, they can be easily detected and have been intensively studied.

Over the span of ~ 5 – 10 Myr, these primordial disks are thought to dissipate (e.g., Strom et al. 1993; Haisch et al. 2001b), based on the observed decline of IR excesses, which originate from warm dust in the inner disk. At older ages, planetesimals and planets comprise the remaining circumstellar population.

Collisions between the planetesimals produce dust grains, which form a “second-generation” disk. The *Infrared Astronomical Satellite (IRAS)* discovered the first example of one such dusty debris disk around Vega (Aumann et al. 1984). Overall, volume-limited far-IR surveys have found that about 15% of main-sequence stars have debris disks (Backman & Paresce 1993; Fajardo-Acosta et al. 1999; Gaidos 1999; Plets & Vynckier 1999; Habing et al. 2001; Laureijs et al. 2002). Such disks have a small amount of dust, little evidence for gas, low characteristic temperatures indicative of grains at

separations of tens of AUs, and a dearth of warm material in their inner regions (e.g., Zuckerman 2001). These attributes make debris disks physically distinct from primordial disks and lead to much weaker observational signatures.

Loss processes for the dust in debris disks act quickly compared to the age of the host stars, and the debris must be continually replenished. Thus, studying debris disks over a wide range of ages can provide insight into their evolution and physical origin. Most studies have focused on stars with ages of ≥ 200 Myr (Habing et al. 1999, 2001; Spangler et al. 2001; Laureijs et al. 2002), a reflection of the fact that most nearby stars are old.

The last few years have seen the discovery of many stars that are both young ($t \approx 10$ – 100 Myr) and close to Earth ($d \lesssim 100$ pc). These stars have largely been found in kinematically associated moving groups by combining accurate space motions with evidence for youthfulness (e.g., Li abundances, X-ray emission, and chromospheric activity). They include the β Pic (≈ 8 – 20 Myr; Zuckerman et al. 2001a) and the Local Association (≈ 20 – 150 Myr; Jeffries 1995; Montes et al. 2001) moving groups. These are important systems for disk studies, as they span an age range that is poorly studied. In addition, this timescale is associated with the assemblage of terrestrial planets in our own solar system (e.g., Wadhwa & Russell 2000), so understanding disk properties at this epoch can illuminate the overall physical context for planet formation.

Given the cold characteristic temperatures of the dust (≈ 40 – 100 K; e.g., Dent et al. 2000), submillimeter continuum observations of debris disks are valuable. The emission is optically thin and hence can be used to estimate the disk masses, with the caveat that large bodies are missed. Previous submillimeter studies have focused on stars already known to have strong IR excesses (Zuckerman & Becklin 1993; Sylvester et al.

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1996, 2001; Holland et al. 1998, 2003; Greaves et al. 1998; Sylvester & Mannings 2000; Holmes et al. 2003), with the most outstanding examples being spatially resolved at submillimeter wavelengths (β Pic, ϵ Eri, Fomalhaut, Vega).

Much less has been done in the way of “blind” searches of stars in a given age range to understand their submillimeter disk frequency and properties. Wyatt et al. (2003) have recently surveyed the secondary stars of binaries with O- and B-type primaries, a sample originally identified by Lindroos (1986). By virtue of the primary stars being on the main sequence, the systems are $\lesssim 100$ Myr old. Previous searches for dust around this sample had relatively poor sensitivity (Gahm et al. 1994; Jewitt 1994; Ray et al. 1995). Of particular interest is the suggestion by Wyatt et al. that some stars may harbor very cold ($\lesssim 40$ K) disks, which would be missed by previous far-IR searches and only detected in the submillimeter. However, one limitation of the Lindroos sample is that most of the systems are fairly distant, which impacts the resulting disk sensitivity of submillimeter observations.

In this paper, we present an $850 \mu\text{m}$ search for cold dust around young stars in the β Pic and Local Association moving groups. Unlike most previous submillimeter studies, our sample was chosen solely on the basis of stellar age. Our sample includes three M dwarfs (GJ 182, 799, and 803) that Barrado y Navascués et al. (1999) identified as the youngest of the nearby M dwarfs, based on the stars’ pre-main sequence location on the color-magnitude diagram. The majority of debris disks from IR studies have been identified around A-, F-, and G-type dwarfs simply because of observational sensitivity bias (e.g., compilations by Greaves & Wyatt 2003 and Zuckerman & Song 2004). The loss processes for circumstellar dust depend on the host star’s luminosity and gravity. Therefore, studying debris disks over a range of stellar masses may also advance our understanding of this phenomenon.

2. OBSERVATIONS

Submillimeter observations were made from 2002 December to 2003 November using the SCUBA bolometer array (Holland et al. 1999) at the 15 m James Clerk Maxwell Telescope (JCMT) on Mauna Kea, Hawaii.⁴ Zenith optical depths ranged from 0.25 to 0.35 at $850 \mu\text{m}$. GJ 803 was first observed on 2003 October 5 and well-detected at $850 \mu\text{m}$; additional $450 \mu\text{m}$ data were obtained on 2003 November 12–13 in dry, stable conditions, with $450 \mu\text{m}$ zenith optical depths ranging from 1.37 to 1.47. We used the photometry mode to maximize observing efficiency, with integration times ranging from 18 to 96 minutes on-source. The data were reduced using Starlink software and standard reduction techniques (see Jenness et al. 2002). Pointing was checked approximately once per hour and was accurate to better than $2''$ rms over each night. Flux calibration factors were derived from observations of Uranus or Neptune when available, but secondary calibrators (CRL 618 or CRL 2688) were also used.

We also observed GJ 803 with SCUBA on 2003 November 12–13 in jiggle-mapping mode. The weather was stable during the observations, with $850 \mu\text{m}$ zenith optical depths ranging from 0.26 to 0.29. Pointing was checked every 1–2 hr using Neptune and was accurate to an rms error of $2''$. Flux calibrations were performed by observing Uranus and Neptune at

the start of each night and were consistent between the two nights to within 10%. A total of 24 individual maps were made at a range of image rotation angles on the array.

GJ 803 was well detected, although the emission is not extended at the signal-to-noise level ($S/N \approx 8$) and angular resolution ($14'' = 140$ AU) of the map. The integrated $850 \mu\text{m}$ flux in the map is consistent with earlier measurements made in photometry mode. The emission is also consistent with being centered on the star, given the pointing uncertainties.

A sensitive search for CO $J = 3-2$ line emission from GJ 803 was conducted with the B3 dual-channel heterodyne receiver and digital autocorrelating spectrometer (DAS) back end on the JCMT on 2003 November 5. Observations were performed in beam-switching mode with a $60''$ throw. Weather conditions were good, and the system temperature ranged from 400 to 450 K over the elevation range of the observations (30° – 40°). A total of 200 minutes of on-source integration was obtained. No signal was detected to a 1σ level of 13 mK per 1 km s^{-1} velocity channel.

3. RESULTS

Table 1 presents our sample and SCUBA photometry. We detected GJ 182 and GJ 803 at $850 \mu\text{m}$, with a possible (formally 2σ) detection of GJ 803 at $450 \mu\text{m}$. GJ 803 is also detected at $60 \mu\text{m}$ in the *IRAS* Faint Source Catalog (FSC; Mathioudakis & Doyle 1991; Song et al. 2002). In addition, we searched for all our targets in the *IRAS* data set using the SCANPI algorithm (Helou et al. 1988) via the Infrared Processing and Analysis Center (IPAC) web site. This search revealed $25 \mu\text{m}$ (color-corrected) fluxes of 65 ± 14 and 157 ± 26 mJy for GJ 182 and GJ 803, respectively, which were not included in the FSC.

For GJ 803, Figure 1 shows the 60 and $850 \mu\text{m}$ excess emission modeled with a single-temperature fit. (The $25 \mu\text{m}$ flux is consistent within 1σ with being purely photospheric in origin.) For the fit, we use a modified blackbody, where the emissivity is constant for $\lambda < 100 \mu\text{m}$ and follows $\lambda^{-\beta}$ for longer wavelengths. The spectral energy distributions (SEDs) of debris disks with more extensive multiwavelength data are well described by this prescription, with $\beta = 1.1$ to 0.5 (Dent et al. 2000). For this range of β , the GJ 803 data are fitted with $T = 35$ – 45 K, with $T = 40 \pm 2$ K for a nominal $\beta = 0.8$, where the uncertainty is the formal 1σ error in the fit. (Fitting a pure blackbody to the data gives 62 ± 5 K.) Note that the *IRAS* nondetection at $100 \mu\text{m}$ is a good match to $\beta = 0.8$ and rules out $\beta \gtrsim 1$. With slightly greater sensitivity, *IRAS* should have detected this source. In addition, the marginal detection at $450 \mu\text{m}$ agrees well with the $\beta = 0.8$ fit but is inconsistent with a pure blackbody. The integrated fractional dust luminosity L_{dust}/L_* is 6.1×10^{-4} for the nominal $\beta = 0.8$, $T = 40$ K model.

For GJ 182, the excess emission at $25 \mu\text{m}$ and $850 \mu\text{m}$ cannot be fit with a single modified blackbody that also satisfies the *IRAS* nondetections. Ignoring the $25 \mu\text{m}$ flux for the moment, a 40 K, $\beta = 0.8$ modified blackbody, like that which fits the GJ 803 SED, is consistent with the 60 , 100 , and $850 \mu\text{m}$ data (Fig. 1). A higher dust temperature would violate the $60 \mu\text{m}$ upper limit. Of course, a lower dust temperature would also be consistent with the $850 \mu\text{m}$ detection. Firmer constraints on the dust temperatures await more sensitive IR/submillimeter measurements.

The mid-IR SEDs are sensitive to the presence of warm dust in the inner disk regions. GJ 182 shows a strong $25 \mu\text{m}$ excess

⁴ The JCMT is operated by the Joint Astronomy Center in Hilo, Hawaii on behalf of the parent organizations Particle Physics and Astronomy Research Council in the United Kingdom, the National Research Council of Canada, and The Netherlands Organization for Scientific Research.

TABLE 1
JCMT/SCUBA OBSERVATIONS

Object	Spectral Type ^a	d^a (pc)	λ (μm)	F_ν^b (mJy)	M_{dust}^b (M_\oplus)
β Pic Group (12^{+8}_{-4} Myr)					
HD 35850 ^c	F7 V	26.8 ± 0.6	850	<4.5	<0.01–0.04
HD 199143	F8 V	47.7 ± 2.3	850	<7.5	<0.05–0.19
HD 358623 (AZ Cap)	K7 Ve	47.7 ± 2.3	850	<6.9	<0.04–0.17
GJ 803 (AU Mic)	M1e	9.94 ± 0.13	850	14.4 ± 1.8	0.011
			450	85 ± 42	
GJ 799 (AT Mic)	M4.5e	10.2 ± 0.5	850	<9.0	<0.003–0.010
Local Association Group (50^{+100}_{-30} Myr)					
EK Dra (HD 129333)	F8	33.9 ± 0.7	850	<5.7	<0.02–0.07
HD 77407	G0	30.1 ± 0.8	850	<4.5	<0.01–0.04
GJ 182 ^d	M0.5 Ve	26.7 ± 1.8	850	4.8 ± 1.2	>0.026
			450	<60	

^a From SIMBAD and the *Hipparcos* catalog.

^b For nondetections, we use the 3σ upper limits on the $850\ \mu\text{m}$ flux and assume $T = 30\text{--}100$ K. For GJ 803, the dust mass is computed assuming $T = 40$ K, derived from the SED fitting. For GJ 182, we assume $T < 40$ K, as described in § 3.

^c Silverstone (2000) reports a possible IR excess based on *ISO* data, but Decin et al. (2000) suggest the excess may arise from extended cirrus, not from the star itself.

^d Montes et al. (2001) show that GJ 182 shares kinematics with the higher mass Local Association stars, which have estimated ages of ≈ 50 Myr based on stellar activity and Li abundances. However, placing GJ 182 on an H-R diagram suggests a younger age of $\approx 10\text{--}20$ Myr (Favata et al. 1998; Barrado y Navascués et al. 1999).

(but see § 4.2). If we add a $150\text{--}200$ K component that emits as a $\beta = 0.8$ modified blackbody, a dust mass of $(4\text{--}12) \times 10^{-5} M_\oplus$ is needed to account for the $25\ \mu\text{m}$ excess, or $0.3\%\text{--}0.9\%$ of the $T = 40$ K dust component that satisfies the $850\ \mu\text{m}$ flux (Fig. 1). For GJ 803, the large mid-IR dip indicates an absence of warm dust in the inner regions. If we add a warm component that emits as a $\beta = 0.8$ modified blackbody, only a very small amount of $150\text{--}200$ K dust is permitted, since the *IRAS* $25\ \mu\text{m}$ flux is consistent with being photospheric. Taking the marginal $25\ \mu\text{m}$ excess at face value gives a warm dust mass of $(3\text{--}9) \times 10^{-6} M_\oplus$, or $0.03\%\text{--}0.08\%$ of the $T = 40$ K dust component.

For the entire sample, we compute dust masses from the $850\ \mu\text{m}$ fluxes in the standard fashion, assuming optically thin emission characterized by a single temperature:

$$M_{\text{dust}} = \frac{F_\nu d^2}{\kappa_\nu B_\nu(T)} \quad (1)$$

where F_ν is the flux density, d is distance, κ_ν is the dust opacity at the observing frequency, and B_ν is the Planck function for a dust temperature T . We adopt a dust opacity of $1.7\ \text{cm}^2\ \text{g}^{-1}$, in agreement with past studies (Zuckerman & Becklin 1993; Greaves et al. 1998; Holland et al. 1998; Sylvester et al.

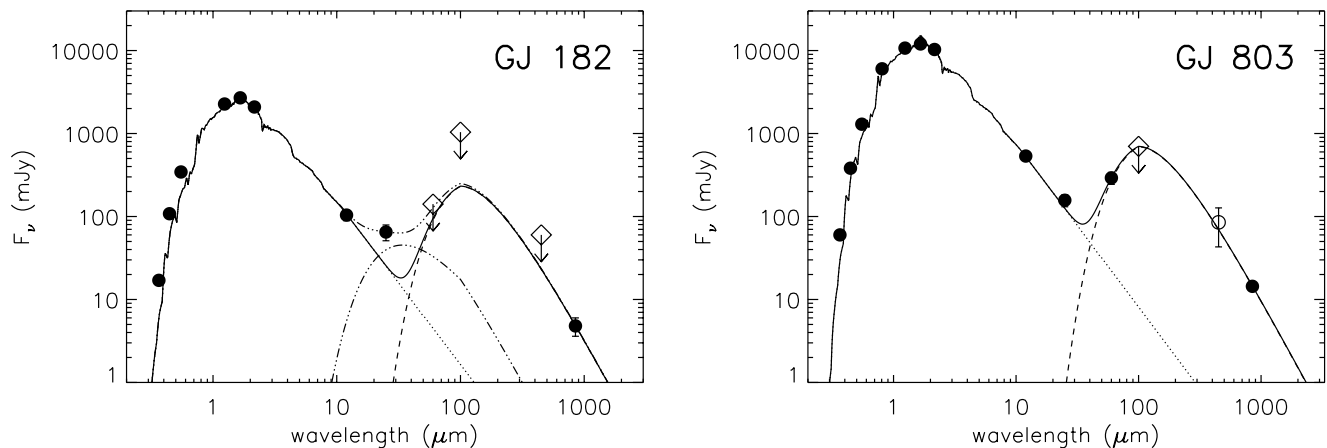


FIG. 1.—Observed SEDs of our stars that were detected by SCUBA. Detections are plotted as filled circles and nondetections as open symbols with arrows. The possible detection of GJ 803 at $450\ \mu\text{m}$ is shown as an open circle. Error bars are shown when they are larger than the symbol size. The data are compared to the sum of a theoretical stellar spectrum (dotted line) and a modified blackbody with $T = 40$ K and $\beta = 0.8$ that fits the $\geq 60\ \mu\text{m}$ detections (dashed line). For GJ 182, we add a 150 K component to account for the possible $25\ \mu\text{m}$ excess (triple-dot-dashed line). See text for details (§§ 3 and 4). The optical data ($0.3\text{--}0.8\ \mu\text{m}$) are from SIMBAD and the *Hipparcos* catalog. The IR data ($1\text{--}100\ \mu\text{m}$) are from the Two Micron All Sky Survey catalog and *IRAS* color-corrected Faint Source Catalog and SCANPI photometry. The submillimeter data (450 and $850\ \mu\text{m}$) are from this paper. The stellar spectra come from the NextGen models by Allard et al. (2000, 2001) and are normalized to the observed K -band magnitude.

2001; Wyatt et al. 2003; Williams et al. 2004; Sheret et al. 2004). This value is on the upper end of the $0.4\text{--}1.7\text{ cm}^2\text{ g}^{-1}$ range discussed by Pollack et al. (1994). For nondetections, we use 3σ upper limits on the $850\text{ }\mu\text{m}$ flux and assume $T = 30\text{--}100\text{ K}$, in agreement with previously detected debris disks (Dent et al. 2000). For GJ 803, we adopt the 40 K temperature from the SED fitting. The calculated dust masses are presented in Table 1.

It is unlikely that our detections are due to background galaxies. Scott et al. (2002) estimate a surface density of about 500 objects deg^{-2} that would be brighter than an $850\text{ }\mu\text{m}$ flux of 6 mJy , about the 3σ sensitivity of our survey. The probability of any background objects to be within an angular distance θ of a target is $1 - \exp(-\pi\theta^2\Sigma)$, where Σ is the surface density of background objects above a specified flux level. We compute the ensemble probability of detecting any background objects in SCUBA's central bolometer for our entire sample of eight objects, accounting for the fact that we mapped GJ 803. This gives a 4% probability that background sources would produce at least one detection in our survey.

The GJ 803 system appears to have very little molecular gas. The CO 3–2 intensity is $-45 \pm 49\text{ mK km s}^{-1}$ integrated over $\pm 7\text{ km s}^{-1}$ about the stellar velocity. (This velocity range corresponds to the maximum orbital speed for a disk viewed at an inclination of 80° with the inner disk edge of 17 AU derived from the SED fit.) Assuming the gas is optically thin and in thermal equilibrium with 40 K dust, the 3σ upper limit implies a CO column density of $6.3 \times 10^{13}\text{ cm}^{-2}$. Temperatures as warm as $\approx 150\text{ K}$, e.g., if the gas were located in the inner few AU, would give inferred CO column densities of a factor of 2 larger. Determining an upper limit on the total gas mass is uncertain, since this is dominated by H_2 , and the H_2 to CO conversion is highly uncertain. CO may freeze out onto grains, making it a poor tracer of the total disk mass. Also, photodissociation may affect the H_2 to CO conversion factor. To attempt to account for photodissociation by the interstellar UV field, we refer to calculations by van Dishoeck & Black (1988), slightly extrapolated to lower CO column densities, and adopt an H_2 to CO abundance ratio of 10^{-7} (cf., normal abundance ratio of 10^{-4}). This results in an H_2 mass of $1.3 M_\oplus$ if we ignore the possibility of CO freezeout. Overall, the gas nondetection limits for GJ 803 are comparable to those for older solar-type stars (Dent et al. 1995; Greaves et al. 2000a) and rule out the possibility of GJ 803 having a gas-rich disk comparable to those detected around younger stars (Zuckerman et al. 1995; Greaves et al. 2000b).

4. DISCUSSION

4.1. Submillimeter Evolution of Debris Disks

Figure 2 summarizes the known debris disk mass estimates based on submillimeter observations from our work and the published literature. Our SCUBA survey adds a significant number of stars of $\approx 10\text{--}50\text{ Myr}$. Our raw $850\text{ }\mu\text{m}$ sensitivity (median rms of 1.9 mJy) is poorer than the Wyatt et al. (2003) survey of Lindroos binaries (median rms of 1.6 mJy). However, given the much closer distances of our targets, our dust mass sensitivity is better by around an order of magnitude. In particular, Figure 2 shows the following:

1. Compared to primordial disks around T Tauri and Herbig Ae/Be stars that have dust masses of $\approx 10\text{--}300 M_\oplus$ (e.g., Osterloh & Beckwith 1995; Natta et al. 2000), the dust masses for the detected young stars are $\sim 10^3$ smaller. This points to very rapid evolution of the circumstellar dust mass within the

first $\sim 10\text{ Myr}$, likely arising from grain removal (e.g., via accretion onto the central star) and/or grain growth into larger bodies, which would greatly diminish the submillimeter/millimeter emission.

2. The upper envelope for disk masses around A stars exceeds that for later type stars. This may reflect the distribution of primordial disk masses; for T Tauri stars a similar trend is seen, where the upper envelope of primordial disk masses is larger for more massive stars (Natta 2004).

3. For stars of $\approx 10\text{--}100\text{ Myr}$ old, nearly all the detected disks at a given spectral type (i.e., stellar mass) are less massive than the upper limits on nondetections. This highlights the fact that existing submillimeter measurements are only sensitive to the most massive of the young debris disks.

4. The apparent correlation of dust mass with stellar age arises from observational bias: the young stars detected in the submillimeter are at larger distances than the old stars. However, the decline in the upper envelope of disk masses with age is probably a real effect. While there are no published submillimeter data in the upper right of Figure 2 (i.e., old stars with massive disks), the existing submillimeter detections of nearby, old stars come from sources known to possess large IR excesses in *IRAS* and *Infrared Space Observatory (ISO)* data. Any old stars with even more massive disks would have been detected by these IR surveys unless the dust was unprecedentedly cold. Based on the data in Figure 2, the masses of submillimeter detected disks evolve roughly as $M_{\text{dust}} \sim t^{-0.5} - t^{-1}$, with a simple unweighted fit giving $M_{\text{dust}} \propto t^{-(0.7 \pm 0.2)}$.

4.2. Debris Disks and Planets around Low-Mass Stars

GJ 182 and GJ 803 are the first M star debris disks detected at submillimeter wavelengths.⁵ The dust lifetime is short compared to the stars' ages (e.g., Backman & Paresce 1993). Therefore, the dust must be replenished by collisions between planetesimals, either in situ or else from regions farther out. Note that for these low-luminosity stars ($0.1\text{--}0.2 L_\odot$), radiation pressure on the grains is negligible compared to the star's gravity, unlike debris disks around A stars, e.g., $\beta\text{ Pic}$, where micron-sized and smaller grains are expelled (Artymowicz 1988).

Our work shows that most debris disks around low-mass stars probably lie at or below the sensitivity limits of the *IRAS* survey and JCMT/SCUBA observations. GJ 182 is just at the practical limit for *IRAS* and SCUBA detection. GJ 803 is the nearest object in our sample and among the nearest known young stars, which makes detection of its very tenuous dust mass feasible. The one undetected M dwarf in our sample, GJ 799, is just as close and young as GJ 803. GJ 799 is known to be a $3''$ binary of comparable magnitude (Mason et al. 2001), which may impact its dust content. In addition, its M4.5 spectral type corresponds to $M_* \approx 0.10\text{--}0.15 M_\odot$, about 3–4 times smaller than that of GJ 182 and GJ 803 (e.g., see models presented in Hillenbrand 1997 and Luhman 1999). Hence, the nondetection may simply be due to insufficient sensitivity; our JCMT data for GJ 799 would have to be ≈ 4 times deeper in order to reach the same M_{dust}/M_* as for the two detected

⁵ The young M3 dwarf Hen 3-600 also shows an IR excess (de La Reza et al. 1989). This binary star is a member of the $\approx 10\text{ Myr}$ old TW Hydrae Association (Zuckerman et al. 2001b). Hen 3-600 has a very large dust optical depth, a significant $10\text{--}20\text{ }\mu\text{m}$ excess (Jayawardhana et al. 1999), and active gas accretion (Muzerolle et al. 2000). These characteristics are distinct from debris disks and indicate that the Hen 3-600 disk is primordial or else in a transitional state.

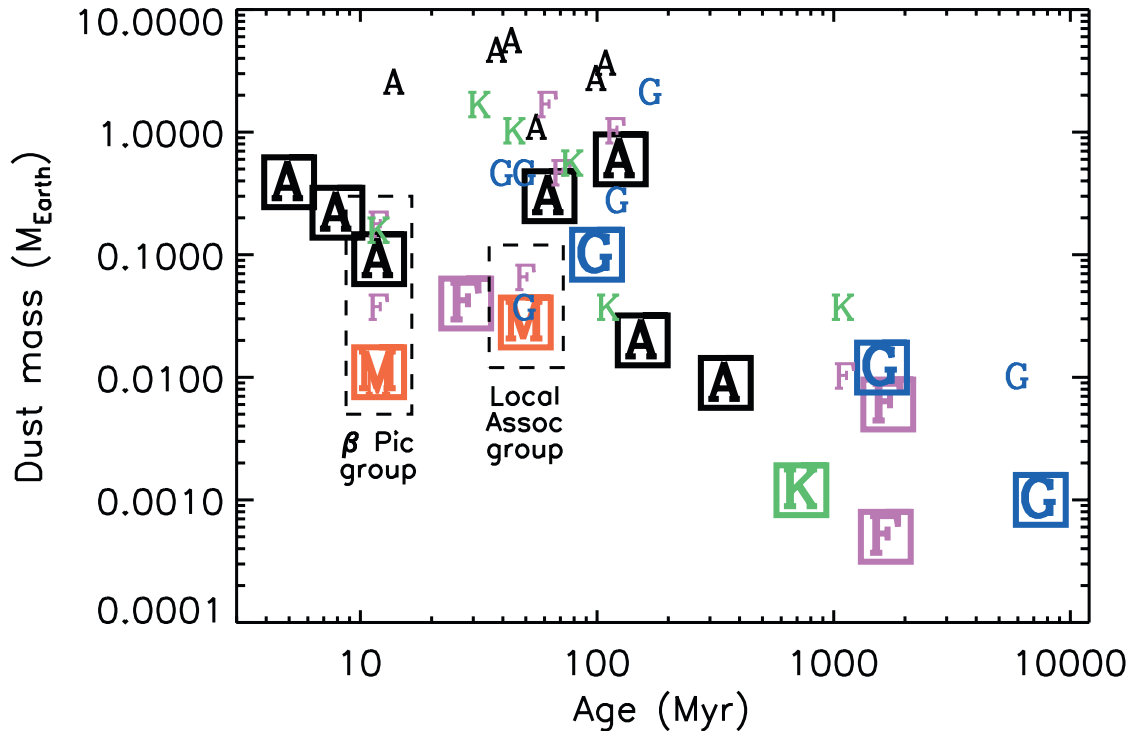


FIG. 2.—Debris disks with dust masses derived from submillimeter data. Detections are shown in boxes, where the letters indicate the spectral type. GJ 182 and GJ 803 are the M dwarfs detected in this paper. Nondetections are shown as smaller letters without boxes, assuming a dust temperature of 30 K. The GJ 799 nondetection is obscured behind the M dwarf detection of GJ 803 at 12 Myr. In addition to this paper, data come from compilations by Wyatt et al. (2003) and Sheret et al. (2004) along with the young G star debris disk from Williams et al. (2004). The quoted age uncertainties range from $\approx 25\%$ to a factor of ≈ 2 .

M dwarfs. Primordial disks are common around low-mass stars (Haisch et al. 2001a) and even brown dwarfs (Liu et al. 2003) at ages of a few Myr, and these systems may very well generate debris disks as they age. The *Spitzer Space Telescope* will offer the requisite sensitivity for detecting and studying such disks.

For GJ 182, the $25\ \mu\text{m}$ excess suggests that dust resides in the inner regions of its disk: 150–200 K blackbody grains would lie at 1–2 AU from the star. Given the sparse sampling of the current SED, one cannot determine if the dust distribution is continuous throughout the disk or if there is a gap in the middle regions of the disk (e.g., from planet formation) between the $25\ \mu\text{m}$ – and $850\ \mu\text{m}$ –emitting dust. However, as shown by Song et al. (2002), single-band excesses near the sensitivity limit of the *IRAS* catalog may not be reliable. Although the SCANPI detection of GJ 182 has $S/N = 4.6$, higher than the $S/N = 2\text{--}4$ objects examined by Song et al., for now we treat the $25\ \mu\text{m}$ excess of GJ 182 as tentative. More detailed SED measurements, in particular with *Spitzer*, would be invaluable to characterize this star’s disk.

The absence of warm dust in the inner regions of GJ 803 (and perhaps GJ 182, if the $25\ \mu\text{m}$ detection is spurious) is naturally explained by an unseen inner companion. The 40 K temperature characteristic of the SED is too cold to be explained by sublimation or melting of icy grains ($\approx 100\text{--}170$ K). Without an inner companion, the effect of Poynting-Robertson drag would cause grains to spiral inward and produce significant mid-IR emission. Such inner holes are typical for debris disk systems. Indeed, the few debris disks with resolved dust emission show morphologies highly suggestive of the dynamical influence of an unseen companion at mean-motion resonances (Holland et al. 1998; Greaves et al. 1998; Ozernoy

et al. 2000; Koerner et al. 2001; Wilner et al. 2002; Quillen & Thorndike 2002; Wahhaj et al. 2003; Kuchner & Holman 2003; Thommes & Lissauer 2003).

The 40 K characteristic temperature from the SED fitting means that blackbody grains around GJ 803 would be at a distance of 17 AU, or an impressively large angular separation of $1''.7$. (For GJ 182, 17 AU corresponds to only $0''.6$.) This provides only a rough guide as to the angular separation of any interior companion, since it is difficult to use the SEDs of debris disks to predict the spatial dust distribution (Sheret et al. 2004). Shallow IR adaptive optics imaging from the Keck Observatory with $0''.05$ resolution shows no stellar or massive brown dwarf companion, suggesting that an even lower mass companion is present. To date, GJ 876 is the only M dwarf known to have extrasolar planets (Marcy et al. 1998, 2001); its two planets have semimajor axes of 0.13 and 0.21 AU.

Thus, the GJ 803 inner disk hole may indicate that planets can exist at larger separations around M dwarfs than known so far and that such planets can form within ≈ 10 Myr. Furthermore, given the very young age of the system, a planetary companion is predicted to have significant thermal emission (e.g., Burrows et al. 1997) and hence might be detectable with deeper adaptive optics imaging. M dwarfs are by far the most numerous type of star, so a better understanding of their planet-bearing frequency is needed for a complete census of the extrasolar planet population.

Follow-up coronagraphic imaging by Kalas et al. (2004) finds that GJ 803 has a large, nearly edge-on disk seen in scattered optical light. This disk is detected from $5''\text{--}21''$ (50–210 AU), although the inner and outer extent are not well constrained. Our submillimeter imaging finds that the $850\ \mu\text{m}$ dust emission is unresolved. This is entirely compatible with

the large optical disk, given that the grains responsible for scattering may be very cold and/or small and hence poor sub-millimeter emitters. Deeper submillimeter mapping is needed for stronger constraints on any extended emission. In general, the proximity and youthfulness of GJ 803, combined with the high submillimeter and optical detectability of its disk, make this system an excellent opportunity to study disk evolution, and perhaps planet formation, in great detail.

The onset and duration of the debris disk phenomenon remain open questions. As discussed by Greaves & Wyatt (2003), the disk masses might slowly decay over time, or perhaps the mass distribution is bimodal, reflecting “on” and “off” states for detectability. The GJ 803 debris disk offers some insight into this issue. Using the *Hipparcos* catalog, which is nearly complete for early-type M dwarfs (M0–M2) within 10 pc, Song et al. (2002) found that GJ 803 was the only object out of a sample of 152 M dwarfs to possess an IR excess in the *IRAS* FSC. This is not due to detection bias, as there are many early-type M dwarfs that have closer distances than GJ 803. GJ 803 is also known to be among the youngest

of the nearby M dwarfs (Barrado y Navascués et al. 1999; Zuckerman et al. 2001a). If the appearance of debris disks was an intermittent event, Song et al. should have detected older M dwarfs with IR excesses at both 60 and 100 μm , if their disks were like the GJ 803 one. Either these older M dwarfs do not have disks, or their disks are much less massive than GJ 803. Therefore, at least for the case of M dwarfs, the GJ 803 system suggests that the time evolution of debris disks is driven primarily by age, with the youngest stars having the largest disk masses.

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